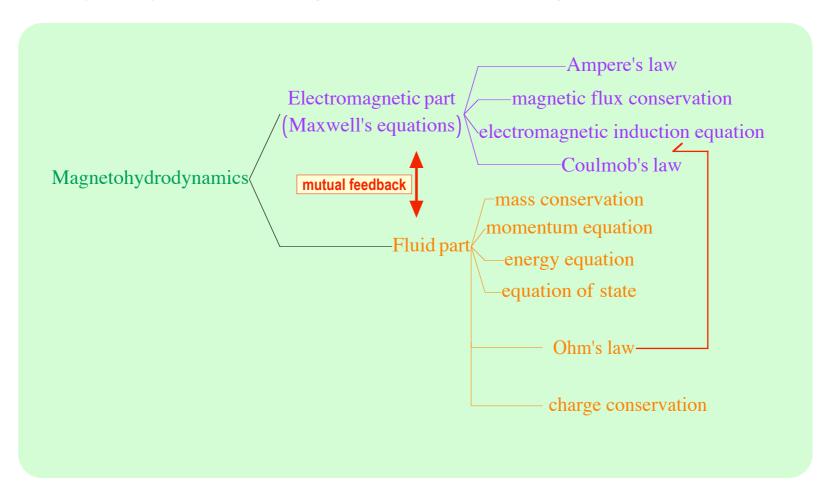
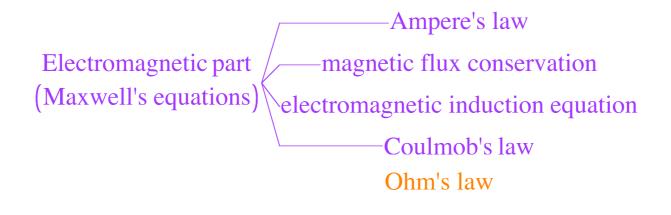
# Magnetohydrodynamics

MHD equations

#### MHD equations...

Fluid (plasma) affects electromagnetic field, while electromagnetic field affects plasma.





(Plasma => Field)

#### **Maxwell's equations** (determine electromagnetic field's state from plasma's state)

 $\nabla \times \boldsymbol{B} = \mu_0 \, \boldsymbol{j} + \frac{1}{c^2} \, \frac{\partial \boldsymbol{E}}{\partial t} \, \dots \, \text{Generalized Ampere's law}$ 

MKS unit

 $\nabla \cdot \mathbf{B} = 0$  ... Conservation of magnetic flux

 $\nabla \times \boldsymbol{E} = -\frac{\partial \boldsymbol{B}}{\partial t}$  ... Electromagnetic induction equation (Faraday's law)

 $\nabla \bullet E = \frac{\rho_c}{\varepsilon_c}$  ... Coulomb's law

 $\epsilon_0 = 8.854 \times 10^{^{-12}}~\mathrm{F~m^{^{-1}}} \qquad \text{... permittivity of free space}$   $\mu_0 = 4\pi \times 10^{^{-7}}~\mathrm{H~m^{^{-1}}} \qquad \text{... magnetic permeability}$   $c = \frac{1}{\sqrt{\epsilon_0~\mu_0}} = 2.998 \times 10^8~\mathrm{m~s^{^{-1}}}~\mathrm{... speed~of~light}$ 

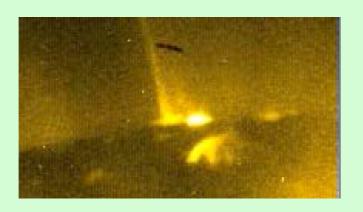
**Unit:** 

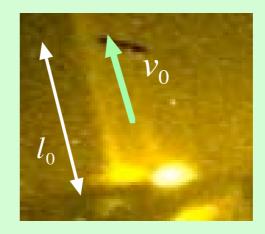
E... V/m, B... T (= 10<sup>4</sup> G), j... A/m<sup>2</sup>

## Typical scales in MHD

MHD phenomena... large-scale, slowly evolving phenomena

#### e.g. solar coronal jet





## Typical scale:

length...  $l_0 \sim$  10^4 km velocity...  $v_0 \sim$  100 km/s

time...  $t_0 = \frac{l_0}{v_0} \sim 100 \text{ s}$ 

 $r_B$  (gyration radius of a proton)  $\sim 1$  m  $t_B$  (gyration period of a proton)  $\sim 10^{-4}$  s in the solar corona

### Approximations in MHD

#### • $v_0 \ll C$ (non-relativistic approximation)

Comparison between electromagnetic terms in the generalized Ampere's law

$$j = \frac{1}{\mu_0} \left( \nabla \times \boldsymbol{B} - \frac{1}{c^2} \frac{\partial \boldsymbol{E}}{\partial t} \right)$$

$$\nabla \times \boldsymbol{E} = -\frac{\partial \boldsymbol{B}}{\partial t} \longrightarrow \frac{E_0}{l_0} \sim \frac{B_0}{l_0}$$

$$\frac{B_0}{l_0} \qquad \frac{1}{c^2} \frac{E_0}{t_0} \longrightarrow \frac{1}{c^2} \frac{E_0}{t_0} \approx \frac{1}{c^2} \frac{1}{t_0} \frac{l_0}{t_0} B_0 = \frac{1}{c^2} \frac{l_0^2}{t_0^2} \frac{B_0}{l_0} = \frac{v_0^2}{c^2} \frac{B_0}{l_0} << \frac{B_0}{l_0}$$

$$\therefore \frac{B_0}{l_0} >> \frac{1}{c^2} \frac{E_0}{t_0} \Rightarrow \text{ we neglect } \frac{1}{c^2} \frac{\partial \boldsymbol{E}}{\partial t}$$

=> Propagation of electromagnetic waves (radiation) is <u>not described</u> by MHD.

$$\nabla \times \mathbf{B} = \mu_0 \, \mathbf{j} + \frac{1}{c^2} \, \frac{\partial \mathbf{E}}{\partial t} \Rightarrow \nabla \times \mathbf{B} = \mu_0 \, \mathbf{j}$$

Ampere's law in MHD (derive magnetic field from current density)

#### • $\rho_c \sim 0$ (local charge neutrality)

$$\nabla \bullet \mathbf{E} = \frac{\rho_c}{\varepsilon_0} \Rightarrow \frac{E_0}{l_0} \sim \frac{e \left(n_+ - n_-\right)}{\varepsilon_0}$$

$$n_+ - n_- \sim \frac{\varepsilon_0 E_0}{e l_0} \approx \frac{\varepsilon_0 \left(\frac{l_0 B}{t_0}\right)}{e l_0} = \frac{\varepsilon_0 v_0 B}{e l_0}$$

#### Condition of local charge neutrality:

 $n_{+} - n_{-} \ll n_{+} + n_{-} \equiv n_{total}$  (total number density)

 $e = 1.6 \times 10^{-19} C$  (electron charge)

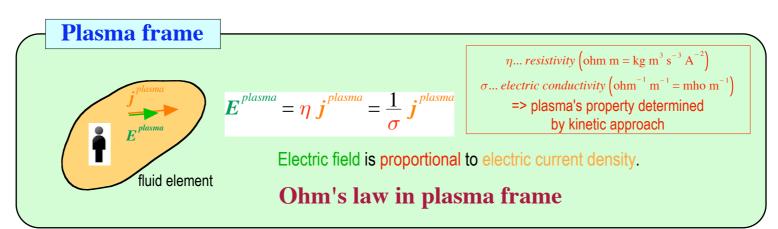
e.g. solar corona:  $n_{total}\sim 10^{14}~{
m m}^{-3},~ v_0\sim 10^5~{
m m/s},~ l_0\sim 10^7~{
m m},~ B\sim 10^{-2}~{
m T}$ 

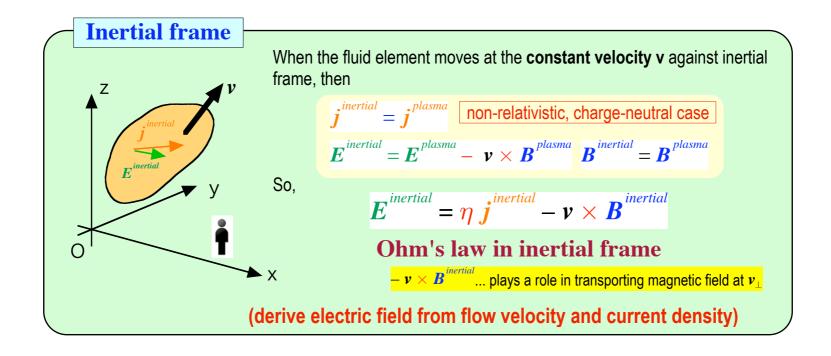
$$6.7 \times 10^7 \frac{v_0 B}{l_0} \ll n_{total} \rightarrow 6.7 \times 10^3 \ll 10^{14} \dots \text{ satisfied}$$

**Coulomb's law in MHD** 
$$\nabla \cdot \mathbf{E} = \frac{\rho_c}{\varepsilon_0} \sim 0$$

(does NOT derive electric field from charge density => Ohm's law)

#### Ohm's law





#### MHD induction equation

... represents the electromagnetic part of MHD

