Coronal mass ejection: global eruption in the corona

intermittent outflow from the Sun

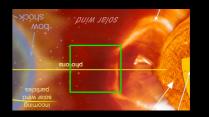


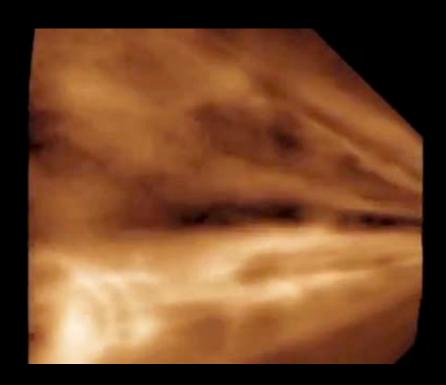


Observed in white light (SOHO)

Solar wind: outflow of a coronal plasma

continuous outflow from the Sun





Solar wind observed by STEREO

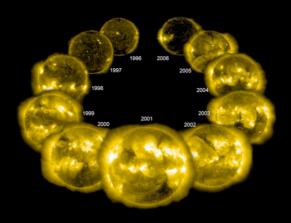
Credits

Animations: NASA Goddard Space Flight Center Data: Craig DeForest, SwRI Research:

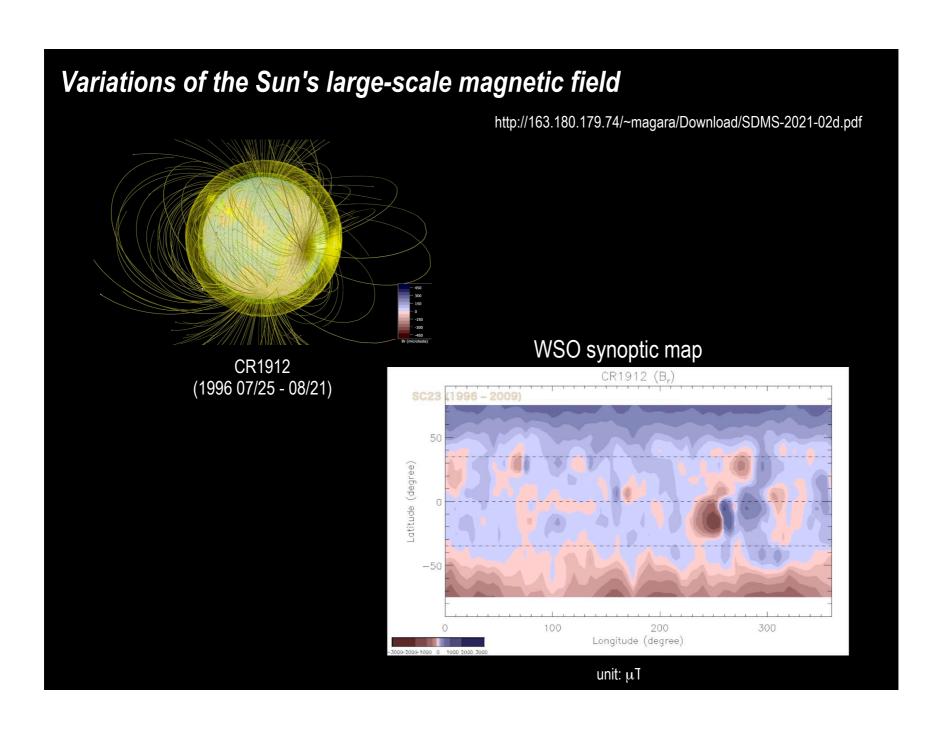
Fading Coronal Structure And The Onset Of Turbulence In The Young Solar Wind C. E. DeForest, W. H. Matthaeus, N. M. Viall, and S. R. Cranmer The Astrophysical Journal, doi:10.3847/0004-637X/828/2/66

Dynamic Sun II

(long time-scale phenomena: ~ years)



Solar cycle: periodic variations of solar global activity Hale's law on solar magnetic cycle Polar field sunspot Beginning of a cycle solar minimum At a deep maximum phase, sunspots are almost randomly distributed around a solar equator. Developing phase Yearly Average Sunspot Number solar maximum Sunspot number 0 Beginning of the next cycle ¹⁹⁴⁰ Year ¹⁹⁶⁰ 1900 2000 solar minimum period... about 11 years



What produces solar dynamic phenomena?





Magnetic fields play important roles in producing solar dynamic phenomena.

- Cyclic amplification of magnetic fields in the solar interior (dynamo)
 - => Produces long-term activity variations known as solar cycle
- Transport of magnetic fields through the convection zone (magnetoconvection)
 - => Lifts magnetic fields against solar gravity via magnetic buoyancy
- Emergence of magnetic fields into the solar atmosphere (flux emergence)
 - => Forms magnetic structure on the Sun (e.g. sunspot, sigmoid, prominence/filament)
- Diffusion of magnetic fields in the solar atmosphere (release of magnetic energy)
 - => Produces explosive phenomena via magnetic reconnection (e.g. flare, jet, coronal heating?)
- Ejection of magnetic fields into the interplanetary space (removal of magnetic fields from the Sun)
 - => Produces outflow/eruptive phenomena (e.g. solar wind, coronal mass ejection)